

COMMISSIONS 27 AND 42 OF THE IAU  
INFORMATION BULLETIN ON VARIABLE STARS

Number 5862

Konkoly Observatory  
Budapest

20 November 2008

*HU ISSN 0374 – 0676*

**DWARF NOVA TRIANGULI 2008 AS A WZ SGE-TYPE OBJECT**

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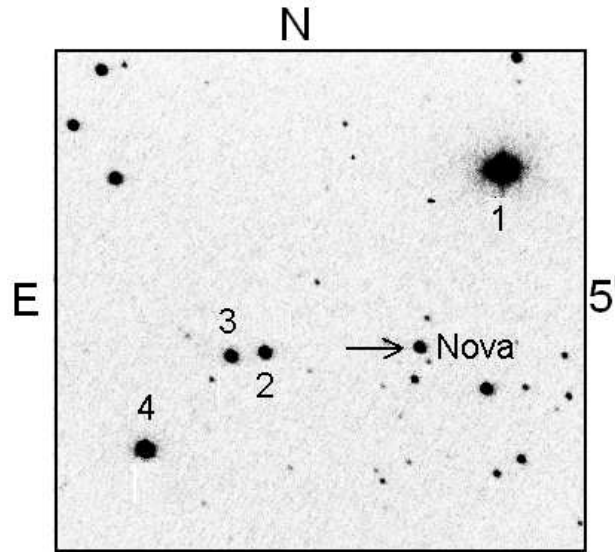
On Oct. 26, 2008, Maehara (2008) reported outburst of the dwarf nova in Triangulum at the coordinates  $\alpha = 02^{\text{h}}38^{\text{m}}39^{\text{s}}.11$ ,  $\delta = +35^{\circ}56'48''.3$  J2000. The brightness of the nova at the time of discovery was  $14^{\text{m}}.12$ , measured with unfiltered CCD. A faint counterpart with  $J = 21.74$  was found in GSC2.3. It indicates large amplitude super-outburst  $\sim 8^{\text{m}}$ , typical for WZ Sge-type dwarf novae. These objects are characterized by the “outburst orbital humps” or “early superhumps”, which appear in the early phase of outburst with the binary orbital period. Nakajima (see Kato, 2008a) did not detect any clear superhumps larger than  $0^{\text{m}}.1$  during the 7.5 hour observations of this dwarf nova before Oct. 28, 2008.

Our  $UBV(RI)_C$  observations of the superoutburst of the dwarf Nova Tri 2008 were obtained with the 0.5 m telescope of the Astronomical Institute of the Slovak Academy of Sciences at Stará Lesná Observatory from October 26, 2008 till November 7, 2008. The SBIG ST10-MXE CCD camera (2184x1472) was used. The part of our CCD frame with the object and comparison stars is presented in Fig. 1.

In this paper we analyse only our  $V$  observations. We used GSC 2336 2105 ( $V = 9^{\text{m}}.95$  according to Hipparcos and Tycho Catalogues (ESA, 1997) and our own measurements) as a comparison star. (No. 1 in the finding chart shown in Fig. 1.) Its constancy was checked against a number of check stars in the field. (No. 2–4 in Fig. 1.)

The light curve of the outburst in  $V$  passband is presented in Fig. 2. The higher resolution of observational runs **A** (early superhumps) and **B** (ordinary superhumps) are shown in Fig. 3.

Fourier period analysis of our CCD observations taken from Oct. 26 till Nov. 1, 2008 (during the first 8 nights of the superoutburst), after trend removal, revealed the presence of small amplitude early superhumps (double-humped variations) with the period  $76.46 \pm 0.5$  minutes ( $0^{\text{d}}0531$ ). This period was used to construct the phase diagram of the residuals. Their mean values with errors are given in Fig. 4. The observations taken from Nov. 2 till Nov. 7, 2008 revealed the presence of  $77.33 \pm 0.2$  minutes ( $0^{\text{d}}0537$ ) superhumps. Their phase diagram is presented in Fig. 5. It is remarkable, that the superhump period of dwarf Nova Tri 2008 is the shortest one among WZ Sge-type objects.



**Figure 1.**  $UBV(RI)_C$  photometric comparison sequence around dwarf Nova Tri 2008.

Similar period values were obtained just recently by Maehara and Ohshima (see Kato, 2008b).

This research made use of the SIMBAD data base, operated by the CDS at Strasbourg, France.

**Acknowledgements.** This study was supported by the VEGA grant No. 7010 of the Slovak Academy of Sciences, SAI scholarship(VIM) and RFBR No. 08-02-01220(SSY).

#### References:

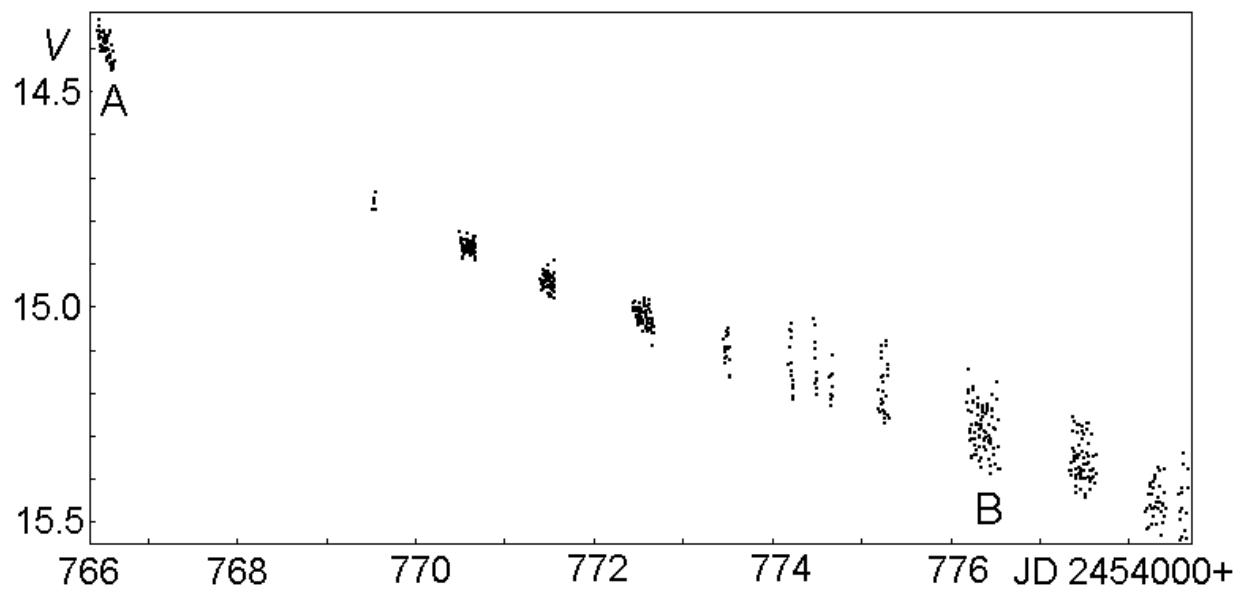
- European Space Agency, 1997, The Hipparcos and Tycho Catalogues, ESA SP-1200  
 Kato, T., 2008a, vsnet-alert 10639  
 Kato, T., 2008b, vsnet-alert 10686  
 Maehara, H., 2008, vsnet-alert 10628

### ERRATUM FOR IBVS 5862

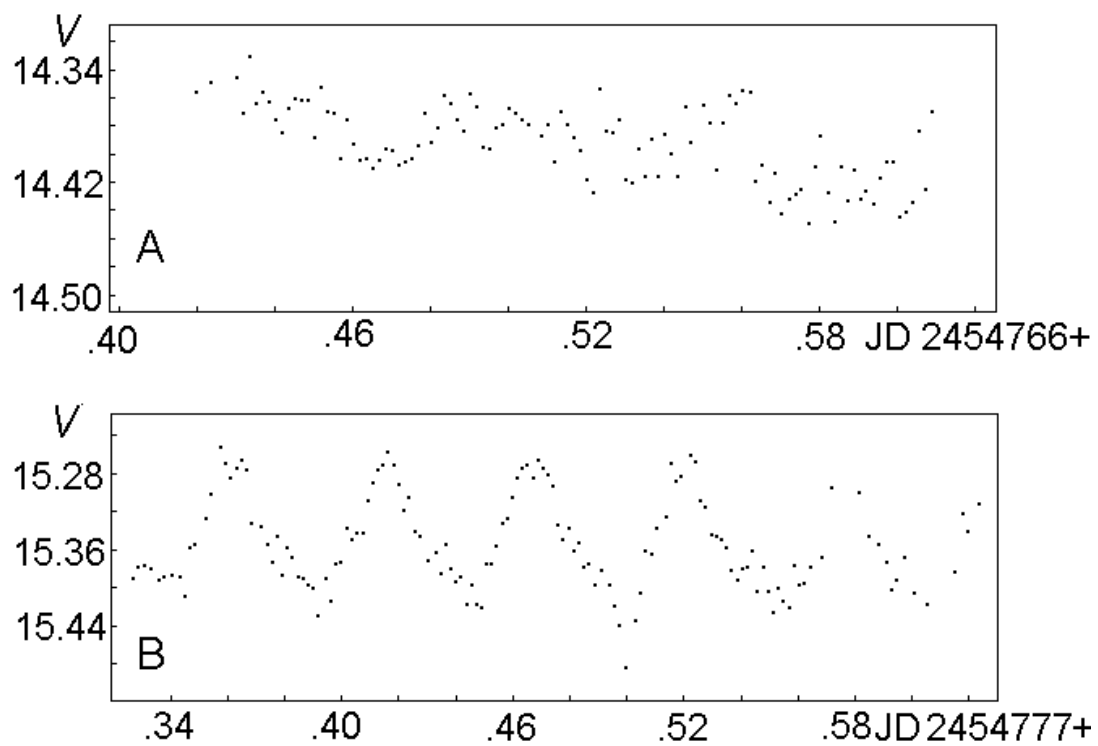
In IBVS 5862, the discovery of the CV was incorrectly attributed to Maehara (2008). The variable was discovered by the Catalina Real Time Transient Survey and designated as CSS081026:023839+355648.

#### Reference:

Cataclysmic Variables detected by CSS, <http://nessi.cacr.caltech.edu/catalina/BrightCV.html>  
 The Editors



**Figure 2.** The light curve of dwarf Nova Tri 2008. The higher resolution of observational runs **A** and **B** are shown in Fig. 3.



**Figure 3.** The early superhumps (**A**) and ordinary superhumps (**B**).

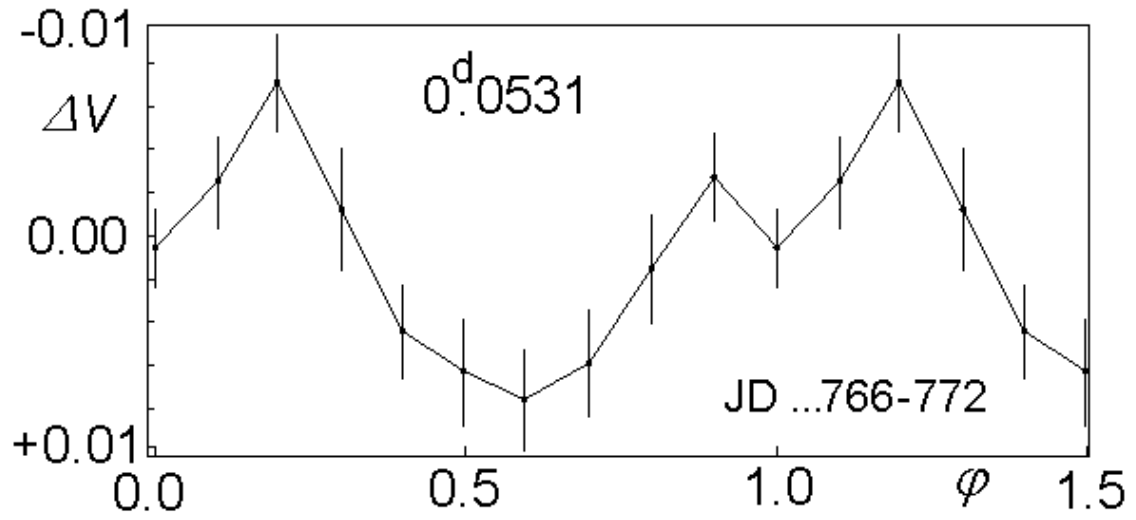


Figure 4. The phase diagram of early superhumps (mean values).

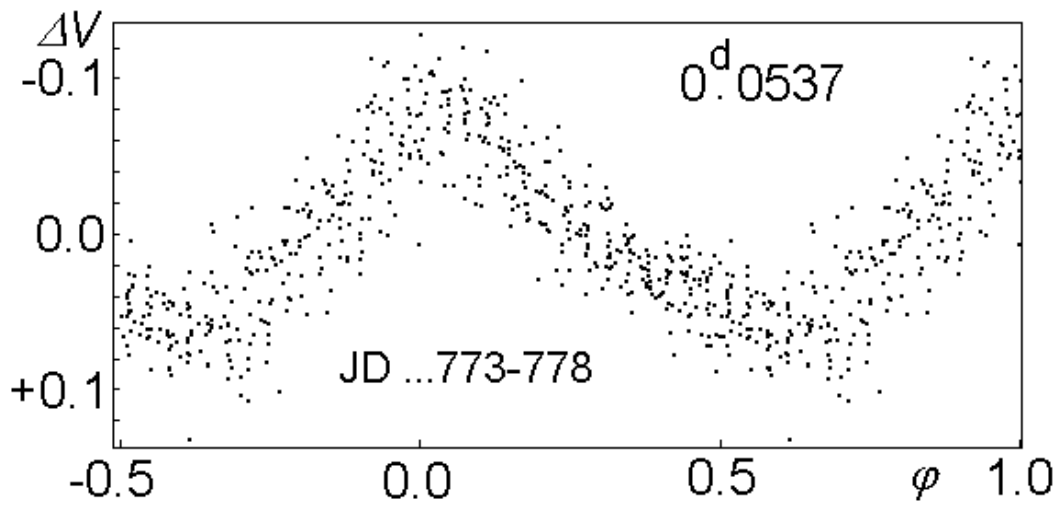


Figure 5. The phase diagram of ordinary superhumps.