## COMMISSIONS 27 AND 42 OF THE IAU INFORMATION BULLETIN ON VARIABLE STARS

Number 4189

Konkoly Observatory Budapest 5 May 1995 *HU ISSN 0374 - 0676* 

## X PERSEI

The Be/X-ray binary X Persei is the optical counterpart of X-ray pulsar 4U 0352+30. Possible orbital period of the system is the 580 day periodicity detected in the radial velocities of the Balmer lines (Hutchings, 1977), in spite of existence of some doubts (Penrod & Vogt, 1985). The star is a known optical variable on time scales from minutes to years. The most comprehensive optical light curve over the past decades is presented by Roche et al. (1993).

X Per has been observed with 60 cm telescope and photon counting photometer of the National Astronomical Observatory Rozhen. HR 1197 (V=6.25, B-V=0.20, U-B=0.14) served as the comparison star. The APR software (Kirov et al., 1993) was used during the data reduction.

Our data are summarized in Table 1. Figure 1 presents the V band light curve over the period 1986-1995. Our observations showed the optical low state that began in the mid-1990 finished in the spring of 1993. Since this time the star has been in optical high state. A clear similarity between the light maximum in 1987-1988 and that in 1993-1994 is visible. The displacement is about 2400 days, i.e. 4 times larger than the supposed orbital period.

The last observations indicate that the optical high state is almost over and X Per is going to a new low state now. May be the star will lose its circumstellar disk again as it was observed in 1990 (Norton et al., 1991; Roche et al., 1993).



Figure 1. V band light curve of X Per. Circles indicate our data. The light curve before JD 2448500 is taken from Roche et al. (1993).

JD 2440000+	V	B-V	U–B	JD 2440000 +	V	B-V	U-B
8545.57	6.81	0.09	-0.68	9341.27	6.47	0.18	-0.78
8620.52	6.76	0.08	-0.64	9343.42	6.47	0.19	-0.78
8712.29	6.78	0.11	-0.73	9357.43	6.44	0.20	-0.78
8734.26	6.76	0.05	-0.68	9366.46	6.44	0.19	-0.78
8841.52	6.75	0.13	-0.66	9572.51	6.34	0.26	-0.73
8842.56	6.76	0.11	-0.67	9573.52	6.33	0.26	-0.73
8905.56	6.78	0.09	-0.70	9597.50	6.26	0.26	-0.73
8916.42	6.78	0.09	-0.70	9598.48	6.27	0.24	-0.73
9028.22	6.78	0.05	-0.66	9715.32	6.38	0.21	-0.75
9046.26	6.76	0.06	-0.66	9748.23	6.45	0.21	-0.81
9195.55	6.45	0.21	-0.75	9772.23	6.48	0.21	-0.83
9197.54	6.42	0.21	-0.76				
9340.28	6.47	0.18	-0.79				

2 able i

Table 1

We gratefully acknowledge Dr. N. A. Tomov for help in observations. This work was partly supported by NSF (F-466/94).

R. K. ZAMANOV NAO Rozhen, POB 136, 4700 Smoljan, Bulgaria

V. I. ZAMANOVA Planetarium, POB 132, 4700 Smoljan, Bulgaria

## References:

Hutchings, J. B., 1977, MNRAS, 181, 619
Kirov, N., Antov, A., Genkov, V., 1993, Compt. Rend. Acad. Bulg. Sci., 44, No. 11, 5
Northon, A. J., Coe, M. J., Estela, A. et al., 1991, MNRAS, 253, 579
Penrod, G., Vogt, S., 1985, ApJ, 299, 653
Roche, P., Coe, M. J., Fabregat, J., McHardy, I. M., Norton, A. J., Percy, J. R., Reglero, V., Reynolds, A., Unger, S. J., 1993, A&A, 270, 122