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**POSSIBLE VARIABILITY OF PLEIADES MEMBER HII 263**

We report here the observation of anomalous dimming seen in relative photometry of the Pleiades cluster member HII 263. As part of a program to monitor young stars in open clusters to derive rotation periods from the observed light variations due to starspots (Prosser *et al.* 1993), the Pleiades cluster member HII 263 was monitored at the Whipple Observatory 48-inch telescope at Mt. Hopkins, Arizona during late October, early November 1992 (Julian dates: 2448922. - 2448933.). CCD photometry of the field containing HII 263 enabled relative V magnitudes between HII 263 and comparison stars of similar brightness to be obtained from each CCD frame. The data will be discussed more fully in a future paper along with the analysis of other stars.

HII 263 is located at RA = 3<sup>h</sup> 41<sup>m</sup> 6.3<sup>s</sup>, DEC = 24° 7' 8''(1950), has  $V \simeq 11.6$ , spectral type G8 (Soderblom *et al.* 1992a) and an observed  $v \sin i$  of 10 km/sec (Soderblom *et al.* 1992b). With such a low  $v \sin i$ , one would predict that HII 263 would possibly have a variability with a period of several days and an amplitude on the order of a few hundredths of a magnitude, similar to that observed among Hyades members (Radick *et al.* 1987). While our data did not reveal any low-amplitude systematic variations, an anomalous dimming of HII 263 was detected on the night of November 3, 1992 (UT). The accompanying figure plots the relative V magnitudes between HII 263 and each of the two comparison stars (A and B) and the relative V magnitudes between the comparison stars over time. The two comparison stars used are HII 309 (=A) and HII 239 (=B), both of which are considered not to be Pleiades cluster members. The observed fluctuation in HII 263 was  $\sim 0.04$  mag, while no variation was detected in the comparison stars. The relative photometry is believed to be accurate to about  $\pm 0.01$  mag. As only one event was observed, no periodicity information for HII 263 could be obtained. Systematic monitoring of HII 263 is recommended in order to determine if the observed variation is periodic, possibly indicating that HII 263 is an eclipsing binary system. If it is an eclipsing system, it must either be a grazing eclipse or the companion must be much smaller than the primary.

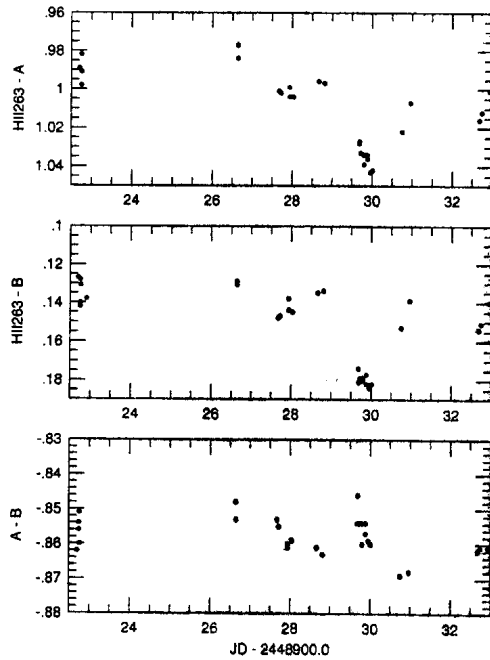


Figure 1

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