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**PHOTOMETRY OF SUSPECTED UU Her-STARS**

Some years ago Sasselov (1984) published a list of suspected UU Her variables based on the spectral classification of Bartaya (1979). Some of these stars were included in a program to determine if they were really variable. A similar work was carried out by Arellano Ferro *et al.* (1989), who concluded that most of the stars in this list were not supergiants, i.e. do not belong to the UU Her-class.

The photometric observations were carried out in Piskéstető with the 50-cm telescope of Konkoly Observatory in the *UBV* system during 1985-1988. The stars were not observed systematically, mainly because of the usual unfavourable spring weather. Table I contains the results of the observations.

**TABLE I**

star	comparison	number of observations	$\Delta V$	$\Delta(B-V)$	$\Delta(U-B)$
HD 51832	HD 50864	4	$-0.734 \pm 0.009$	$-0.037 \pm 0.003$	$+0.157 \pm 0.007$
HD 56664	BD+61°968	2	$+0.746 \pm 0.006$	$+0.071 \pm 0.004$	$+0.035 \pm 0.001$
HD 58965	HD 59506	6	$+1.258 \pm 0.017$	$-0.650 \pm 0.005$	$-0.972 \pm 0.009$
HD 58966	HD 59506	6	$+0.873 \pm 0.004$	$-0.665 \pm 0.009$	$-0.984 \pm 0.004$
HD 102102	HD 101549	4	$+0.117 \pm 0.002$	$+0.152 \pm 0.004$	$+0.008 \pm 0.012$
HD 102570	HD 102482	4	$-0.313 \pm 0.006$	$-0.539 \pm 0.004$	$-0.598 \pm 0.012$
HD 130667	BD+47°2178	6	$-0.967 \pm 0.007$	$-0.091 \pm 0.005$	$+0.066 \pm 0.012$

It is clear that these stars are not variables. HD 58965 may be the only exception, but if the larger scatter is due to variability, it probably is not of the UU Her-type (the amplitude is smaller than 0.05 mag). There are only two measurements of HD 56664 with differences of 0.008 mag in *V*, 0.005 mag in *B-V* and 0.001 in *U-B*, but there is a three year long gap between them. Thus the *UBV* photometry of these stars supports the conclusion of Arellano Ferro *et al.* (1989) that these stars are definitely not UU Her-type variables.

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