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VARIABILITY IN THE HIGH LATITUDE Be STAR JL212 (=CPD-56°154)

Star number 212 in the survey by Jaidee & Lynga (1969) has Strömgren photometry by Kilkenny (1984) which indicates it to be of spectral type near B2. An analysis by Keenan, Dufton & McKeith (1982) found JL212 to have normal abundances for a population Istar although it apparently lies at z = -3.2 kpc from the galac-They also noted that differences between surface tic plane. gravity determinations using spectroscopic measurements of HE and photometry of the Hß line suggested that JL212 might be an emission line star. Subsequently, spectroscopy by Kilkenny & Muller (1988) showed that JL212 did indeed have Hβ in emission. et al (1982) found ( $T_{eff}$ , log g, V sin i) = (19200, 3.7, 219) from high dispersion spectroscopy; using moderate dispersion (30 A/mm) Kilkenny (1989) obtains (17500, 3.5, 225) with a distance from the galactic plane z = -3.3 kpc, which is in good agreement with the Keenan et al (1982) result and is also extremely unusual for a Be star.

A very similar star, SB357 (=CD-37°316) at z = -6 kpc has recently been shown by Kilkenny (1988) to be a variable star with no obvious short-term (~ hours) variation but clear variability on a time scale of a few days or longer. It therefore seemed worthwhile to examine the SAAO 'archive' photometry of JL212 in a search for possible variability. The uvby and UBV(RI) $_{\rm C}$  data are summarised in Tables 1 and 2 where the second row of figures for each date are the standard deviations of the mean values in units of 0.001 mag. All data were corrected to observations of the nearby star CPD-56°153 for which we find:

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s.d
                       s.d
                                     y = 9.805 \pm 0.004
 ٧
          = 9.802 \pm 0.009 44
                                   (b-y) = +0.546 \pm 0.005  15
         = +0.895 \pm 0.005 44
(B-V)
                                     m_1 = +0.258 \pm 0.007 15
          = +0.494 \pm 0.007 44
(U-B)
                                     c_1 = +0.439 \pm 0.014  15
          = +0.480 \pm 0.004  17
(V-R)
(V-I)<sub>c</sub>
          = +0.951 \pm 0.005 17
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Table 1 Strömgren photometry of JL212 (=CPD-56°154)

n	c <sub>1</sub>	$m_1$	(b-y)	٧	НЭО
2	+0.291	+0.076 4	-0.062 1	10.320 2	2444109.51
5	+0.277 8	+0.089 5	-0.067 3	10.345 5	4119.54
2	+0.281 7	+0.080 2	-0.065 1	10.340 2	4120.58
1	+0.261	+0.087	-0.060	10.313	5251.43
1	+0.278	+0.085	-0.065	10.318	5252.45

Although CPD-56°153 is not an ideal comparison star for JL212, being much redder (early K by colour) it is of similar brightness and is only ~ arcminutes away from JL212. CPD-56°153 was always reduced to the usual 'all-sky' standards for UBV(RI)\_C (Menzies, Banfield & Laing 1980) and uvby photometry (Crawford & Barnes 1970) and appears to be of constant brightness.

For JL212 there is only one reasonably long sequence of monitoring, that on HJD 2447005, which shows no variation bigger than about 0.01 mag over  $\sim$ 2.5 hours. From night to night comparison, it seems quite clear that JL212 is variable with an observed range of  $\sim$ 0.04 mag. The data obtained are from 1979, 1981 and 1982 (uvby) and 1987 (UBVRI) and show no sign of the much larger scale variations ( $\sim$ 0.2 mag) seen in SB357.

Table 2  $UBV(RI)_{C}$  photometry of JL212

n	(V-I)	(V-R)	(U-B)	(B-V)	٧	нэр
28			-0.646 3	-0.158 3	10.329	2447005.51
6			-0.653 3	-0.155 3	10.330	7006.62
1	-0.163	-0.070	-0.656	-0.160	10.341	7007.62
11			-0.652 4	-0.158 4	10.323	7009.61
7			-0.654 4	-0.165 5	10.326 3	7011.58
4			-0.651 3	-0.154 4	10.343 2	7013.55
18			-0.651 3	-0.157 6	10.321 5	7021.62
2	-0.142 4	-0.061 10	-0.661 4	-0.152 1	10.316 4	7023.59
8	-0.145 9	-0.063 5	-0.656 2	-0.158 5	10.351 2	7025.62
2	-0.150 9	-0.065 8	-0.650 l	-0.155 1	10.356 2	7068.49
2	-0.149 9	-0.071 1	-0.647 1	-0.157 1	10.347 1	7070.52
2	-0.156 4	-0.068 2	-0.646 1	-0.164 2	10.351	7071.47
2	-0.139 8	_	-0.654 2		10.309	7078.43

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