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EF PEGASI

Tsesevich et al. (1979) gave a report on this U Geminorum type object, which is separated by only 5" from a 15.5 magnitude star.

In order to search for further brightenings, I performed magnitude estimates on about 900 sky patrol plates covering the time interval 1928-1986. Because of the low limiting magnitude of the plates, only the brightness maxima, which are given in Table I could be observed (about $11^m.6 \dots 12^m.0$, near the brightness of the comparison star "c" in the paper of Tsesevich et al.). The maximum at J.D. 243 6071 given there was published by Tsesevich et al. (1979). As the maxima recur at rather regular intervals, they can be described with the following elements:

$$\text{Max} = 242\ 6335 + 162^d.5.E.$$

A similar cycle length was conjectured by Hoffmeister (1935).

Table I

J.D.	E	O-C
242 6333	0	- 2 ^d
6987.88	4	+ 2
243 0547	26	-13
5781	58	+21
6071	60	-14
7906-11	71	+36
8680	76	- 5
9672	82	+12
244 5165	116	-20

The separation of EF Peg from the 5" comparison is, of course, not possible on our plates. Inspecting the figure on page 5 of the paper by Tsesevich et al. (1979), one may suppose that EF Peg belongs to the SU UMa subclass of dwarf novae. This question, however, could not be answered because brightness rises and declines could not be followed up on Sonneberg plates.

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References:

- Hoffmeister, C., 1935, *Astron. Nachr.* 255, 407.
Tsesevich, V.P., Goranskij, V.P., Samus, N.N., Shugarov, S.Yu., 1979, *Astron. Tsirk. No.* 1043.