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NEW FLARE STAR SVS 2559 PERSEI

The variability of SVS 2559 was discovered by Kurochkin (Astron.Tsirk. No. 1325, p. 5; 1984), who already suspected that the object might be a flare star. His plate material, however, accidentally did not allow to trace completely one of the eruptions.

On 343 suitable Sonneberg 40 cm astrograph exposures of the years 1964 to 1983 one conspicuous flare could be detected confirming the UV Ceti type. It is recorded by two of five plates of a night series of 1971 Oct. 19 as follows:

J.D.	244 1244.283	16.7 ^m pg
	.346	16.7
	.401	15.1
	.457	16.4
	.557	16.7

As the exposure times of the plates amount to 80 minutes each, the observed range of the eruption is flattened, by an unknown amount.

The distribution of brightness data of our material very nearly resembles that given by Kurochkin (l.c.). We should note that the star is never invisible on good plates. Therefore the "faint tail" of the distribution is supposed not to be caused by observational bias. The comparison stars used are those of Kurochkin.

This variable is (besides QZ Persei, of similar brightness) one of the two UV Ceti stars to lie in the Sonneberg 100 square degrees field centred at ρ Persei.

W. WENZEL
Sternwarte Sonneberg
Zentralinstitut für Astrophysik
der Akademie der Wissenschaften
der DDR