COMMISSION 27 OF THE I. A. U.
INFORMATION BULLETIN ON VARIABLE STARS

Number 2162

Konkoly Observatory
Budapest
1982 June 15

HU ISSN 0374-0676

IMPORTANCE OF UBV OBSERVATIONS OF AU Mon DURING JANUARY-FEBRUARY 1983

The mathematical analysis of the photometric peculiarities of AU Mon (Lorenzi, 1980a,b) has allowed to separate the variable light curve into two overlapped periodic phenomena, one of the variations due to eclipses, the other to intrinsic variation, with the ephemerides:

$$M(E) : 2442801.3752 \pm {}^{51} + 11.1130371 \pm {}^{68} E \quad (1)$$

$$Max(E) : 2443105 + {}^{4}411 E \quad (2).$$

In order to describe univocally the previous brightness behaviour of AU Mon, the light surface device (Lorenzi, 1980b) has been introduced.

At present we would emphasize the importance of future UBV observations of AU Mon during January-February 1983. In fact the ephemeris (2) involves the occuring of a brightness minimum of the intrinsic variation just in the middle of this period, which, besides, is the best epoch for observing AU Mon. A photometric snapshot of the eclipsing variation relating to such period should allow a geometrical solution of the system to compare both with the average one already computed (Lorenzi, 1982) and with another solution referring to the maximum of the intrinsic variation.

L. LORENZI

Casella postale 24
10025 Pino Torinese (To)
Italy

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