

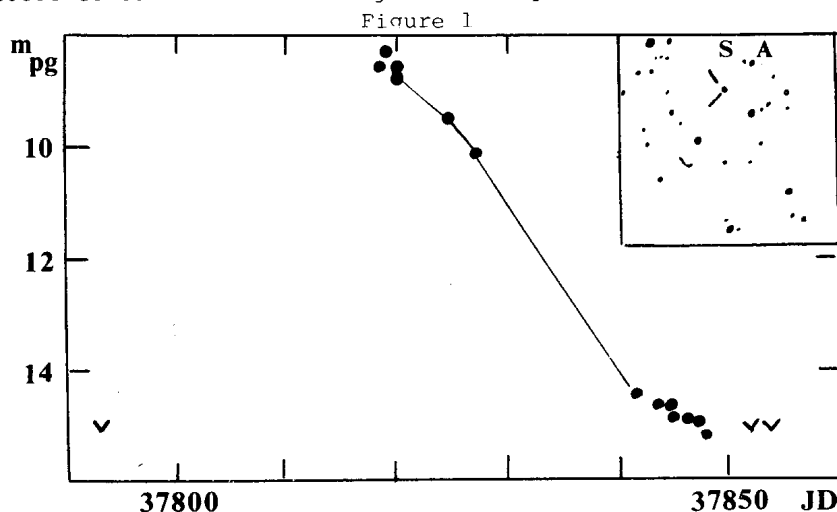
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INFORMATION BULLETIN ON VARIABLE STARS

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Konkoly Observatory  
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NOVA SAGITTARII 1962

With the Rodman blink microscope I discovered a probable nova on two plates taken by Dr. Dirrit Hoffleit with the 7.5-inch Cook triplet at the Maria Mitchell Observatory. The magnitude of the nova was about 8.4 pg on the plate taken on JD 37818, and it was fainter than the plate limit (about 15<sup>m</sup>.0) on JD 37929. Plates taken on the intervening days yielded the light curve in Figure 1. Unfortunately there were no previous plates closer in time to JD 37818 than 37795, so that the exact date of outburst cannot be determined from these data. On JD 37795 the star was fainter than 15<sup>m</sup>.2, the approximate plate limit, and it returned to this magnitude within a period of between one and two months after outburst. That the star is a nova is concluded from the shape of the light curve and the fact that it did not reappear on the Nantucket plates covering JD 36000 to the present, nor on Harvard patrol plates of the RB-RH series between JD 26000 and 33000, which reach an average limiting magnitude of 14.5. Some 700 plates were examined. About 70 scattered Harvard A plates from JD 20000 to 30000 showed no object as bright as 15<sup>m</sup>.0.\*

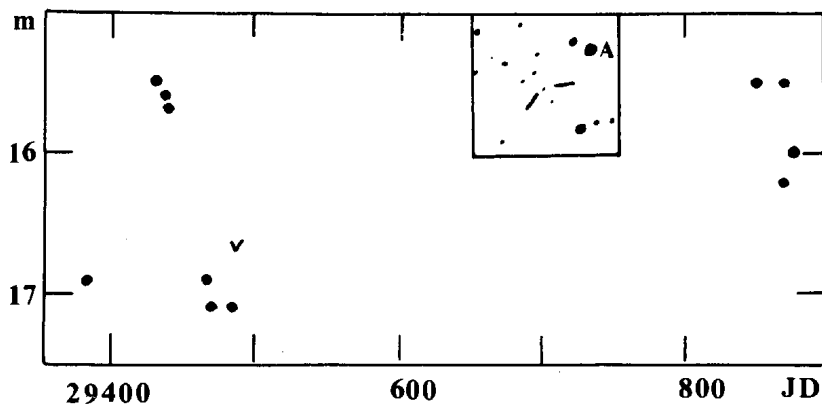


Photographic magnitudes for the brighter points were estimated using the Smithsonian Astrophysical Observatory Star Catalog magnitu-

des of surrounding stars corrected from visual to photographic. For the fainter points a nearby Harvard sequence was used. The approximate position, interpolated from the Smithsonian Catalog position of four nearby stars, is  $18^{\text{h}}24^{\text{m}}36^{\text{s}}, -24^{\circ}04'.8$  (1900).

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Figure 2



\* ADDENDUM - The Harvard A-plates (24-inch Bruce refractor, scale 60"/mm in contrast to 248"/mm for the Nantucket plates) show a faint variable star in the position of the nova. It varies from  $15.5^{\text{m}}$  to  $17.0^{\text{m}}$ . This star is also weakly visible on a few plates of the Harvard B and MF series, but there are too few positive observations to ascertain whether or not the variation is periodic. Figure 2 shows the most definitive run of observations, together with an enlarged diagram showing fainter stars than the insert in Figure 1. The star marked A in both diagrams is CoD -24<sup>o</sup>14410.

DORRIT HOFFLEIT