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A mean light-curve of this RW Aur-type star has been constructed from Tsessevich's observations performed on Harvard Sky Patrol plates and from Moscow and Odessa plates (Fig.1). There is one big and slow fluctuation of brightness with a very long cycle, besides there are smaller superposed oscillations. We have determined the length of individual cycles and plotted them in Fig.2. It is clear that the cycles get shorter as the mean brightness of the star increases. The cycles vary within the limits of 500 to 2500 days.

There are also rapid fluctuations. When the star gets brighter we observe short Algol-like weakenings (Fig.3). If the brightness of the star gets fainter, its behaviour changes: we observe rapid flares (Fig.4).

The comparison of photographic and visual observations shows that the visual magnitude is not a simple function of the colour-index. The relationships between the two quantities for different time intervals are given in Fig.5 and in the following Table:

Interval	Relationship between V and CI
I 2425209-2427358	$m_V = 11.89 - 1.91 \text{ CI}$
II 2427407-2428170	$m_V = 10.72 - 1.51 \text{ CI}$
III 2428441-2429326	$m_V = 10.58 - 1.01 \text{ CI}$
IV 2437254-2438734	$m_V = 12.06 - 1.59 \text{ CI}$
V 2439022-2439480	$m_V = 15.57 - 4.95 \text{ CI}$

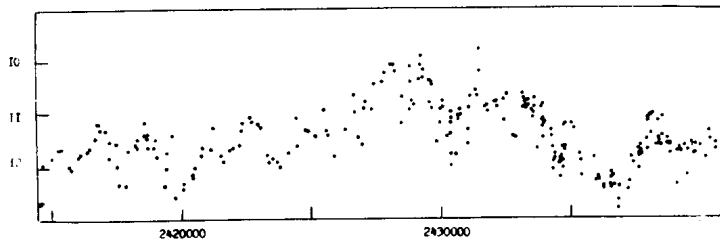


Fig. 1.
General mean light-curve of RY Tauri

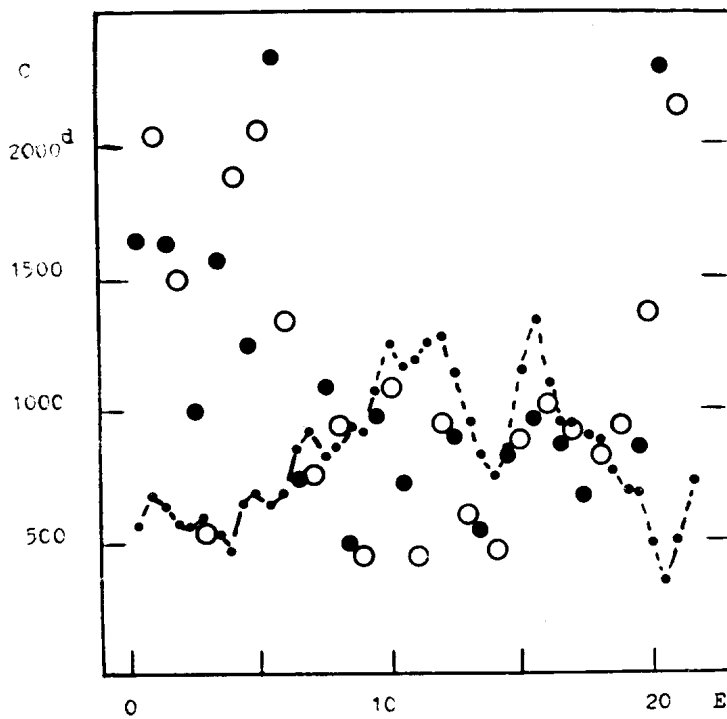


Fig. 2.
C: length of cycle, E: serial number of cycle.
Dotted line: curve of mean brightness.

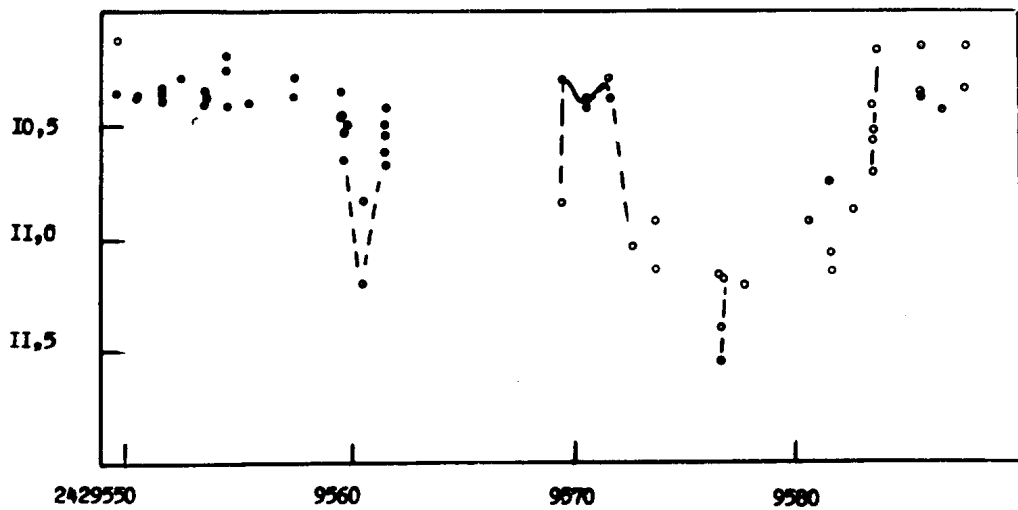


Fig. 3

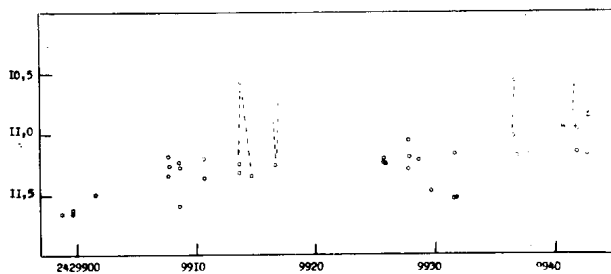


Fig. 4

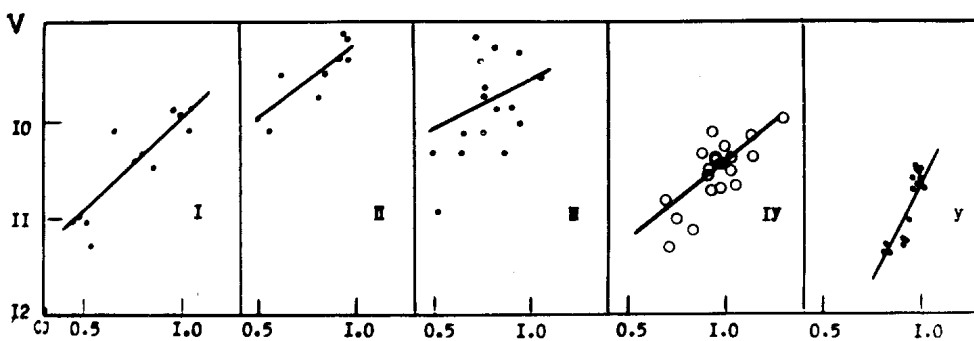


Fig. 5