

COMMISSION 27 OF THE I. A. U.
 INFORMATION BULLETIN ON VARIABLE STARS

NUMBER 351

Konkoly Observatory
 Budapest
 1969 May 30

THE PERIOD OF EV Sct

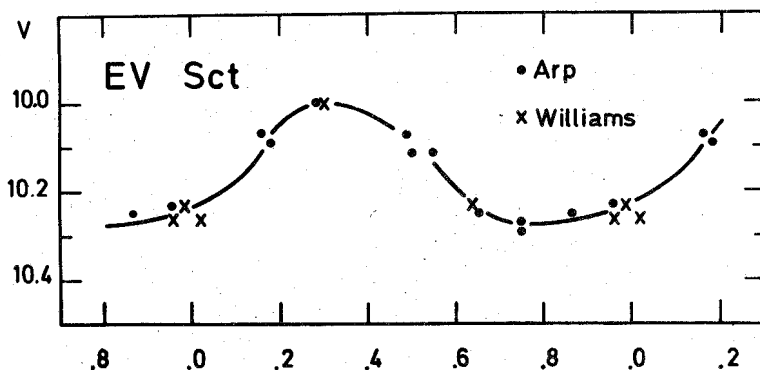
Only 12 cepheids in galactic clusters are available for the calibration of the period-luminosity-colour relation (1). It is therefore surprising that the data for one of them, EV Sct, are uncomplete. First, the reality of the cluster NGC 6664, of which it is supposed to be a member, is still questionable as pointed out by Bakos (2) and J. Stock (cited by Becker, 3). Secondly, the light curve and the period of EV Sct as derived by Arp (4) is based on only 11 photoelectric observations.

It is the aim of this note to improve at least the latter situation. Williams (5) has observed EV Sct by narrow-band photometry in five nights and has derived V-magnitudes of the UBV-system. There is no indication that systematic errors are introduced by this magnitude transformation. Using these additional observations the period can be improved to $3^d 0910 \pm 0^m 0002$. This is sufficiently accurate that the much earlier observations by Bakos and Oosterhoff (cited by Bakos, 2) can be unambiguously combined with the later observations. One finds then the elements:

$$J.D.(Max) = 2\ 436\ 120.998 + 3^d 09097 . E$$

$\begin{matrix} +15 & +6 \end{matrix}$

As seen in the figure this period brings the observations



by Arp (4) and Williams (5) to very good agreement. There is therefore no need to adjust the other photometric elements published by Arp (4).

Astronomisches Institut der Universität Basel
May 18th, 1969

G.A. TAMMANN

References

1. Sandage, A., and Tammann, G.A. 1969, Ap.J. 157 (in press)
2. Bakos, G.A. 1950, Leiden Ann. 20, 177
3. Becker, W. 1963, Zeitschr. f. Astroph. 64, 77
4. Arp, H.C. 1958, Ap.J. 128, 166
5. Williams, J.A. 1966, A.J. 71, 615