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FLARE ACTIVITY OF YZ CMi

The flare star YZ CMi ( $M_v = 11^m.6$ ,  $RA = 7^h 42^m.9$ ,  $D = 3^\circ 39'$  (1968)) has been monitored at the Boyden Observatory during the dark moon-periods between 1968, Jan. 23/Febr. 5 and 1968, Febr. 21/March 7. The star was observed in Johnson's spectral region B with the 16" Harvard Nasa Cassegrain reflector; the star-intensity was measured against the sky-back-ground only, without observing comparison stars. A total period of  $67^h 15^m$  was covered. The table contains information about the time intervals of coverage and all 10 flares occurred. The figure shows the three most prominent flares and one minor, but complex event with more detail.

Date (1968)	Observing time (U.T.)	Cover- age	Time of flare (max.) (U.T.)	Flare N <sup>o</sup>	$\Delta m$	Duration of flare			
						before max.	after max.	Total	
1. 23/24	19 <sup>h</sup> 50-22 <sup>h</sup> 42	4 <sup>h</sup> 16 <sup>m</sup>	21 <sup>h</sup> 29 <sup>m</sup> .6	1	1 <sup>m</sup> .65	0 <sup>m</sup> .8	≈21 <sup>m</sup>	≈22 <sup>m</sup>	
	23 00- 0 24		22 09.0	2	0.39	0.4	3.3	3.7	
			23 23.6	3	0.43	0.2	0.4	0.6	
	24/25	19 32-20 58	1 26						
	25/26	18 37-20 26							
		23 31- 1 02	3 20						
28/29	18 20- 1 03	6 27	0 29.7	4	0.36:	0.8	2.5	3.7	
	29/30	20 20-23 24	20 29.7	5	0.50	1.0:	1.5	2.5	
			21 39.7	6	1.50	0.4	≈24	≈24	
30/31	18 55-20 14	1 19							
1.31/2.1	21 41- 0 20	2 39							
2. 1/ 2	18 25- 0 41	6 06							
	2/ 3	18 24-23 08	4 36	19 14.3	7	0.35:	0.3	0.5	0.8
	5/ 6	18 25-19 58							
		21 33-22 30	2 30						
TOTAL:		35 <sup>h</sup> 43 <sup>m</sup>							
2. 21/22	18 25- 0 05	4 40	20 40.8	8	0.70	0.8	1.7	2.5	
	22/23	17 56-18 39							
		21 37-23 45	2 51						
	25/26	18 12- 0 05	5 53	18 41.3	9	0.40	0.2	0.9	1.1
				21 50.4	10	1.20	0.8	≈11	≈12
27/28	18 00-23 58	5 58							
28/29	18 00-23 50	5 50							
2.29/3.1	18 05-22 00	3 55							
3. 7/ 8	21 35- 0 00	2 25							
TOTAL:		31 <sup>h</sup> 32 <sup>m</sup>	: uncertain values						

The other 6 flares show the normal simple appearance: a steep increase before and an exponential decrease of intensity after reaching the maximum. The measurements of flare Nr. 7 and - to a smaller degree - also of Nr. 4 are not very reliable; these flares are to be checked with other observations.

